

Detection of $^{13}\text{C}17\text{O}$ line emission in the disk around the young outbursting protostar V883 Ori by ALMA observations

S. Notsu,^{1,2} T. Tsukagoshi,³ H. Nomura,⁴ T. Hirota,⁴ M. Honda,⁵ E. Akiyama,⁶ A. S. Booth,⁷ C. Walsh,⁸ T. J. Millar,⁹ S. Lee,¹⁰ J. Lee,¹¹ Y. Yamato,¹ Y. Aikawa,¹ Y. Okoda,² N. Sakai,²

¹*Department of Astronomy, The University of Tokyo, Japan* ²*Star and Planet Formation Laboratory, RIKEN, Japan* ³*Ashikaga University, Japan* ⁴*National Astronomical Observatory Japan, Japan* ⁵*Okayama University of Science, Japan* ⁶*Niigata Institute of Technology, Japan* ⁷*Harvard-Smithsonian Center for Astrophysics, USA* ⁸*University of Leeds, UK* ⁹*Queen's University Belfast, UK* ¹⁰*Korea Astronomy and Space Science Institute, Korea* ¹¹*Seoul National University, Korea*

Understanding the spatial distribution of gas and mass evolution in protoplanetary disks are essential for elucidating the formation and evolution processes of planetary systems. Therefore, a lot of statistical observations of the disk gas distribution using $^{13}\text{C}16\text{O}$ and $^{12}\text{C}18\text{O}$ emission lines have been carried out (e.g., [1]). However, if these emission lines are optically thick, the disk gas mass will be underestimated. Here, by using the $^{13}\text{C}17\text{O}$ emission line, which is the rarest and thus most optically thin among the stable carbon monoxide isotope species, it becomes possible to measure the gas mass with higher precision, including disk midplane. In recent ALMA observations, the $^{13}\text{C}17\text{O}$ emission line was detected for the first time in two protoplanetary disks (HD 163296: [2]; HL Tau: [3]), and the estimated disk masses are about 2-10 times heavier than those estimated from previous $^{12}\text{C}18\text{O}$ emission line observations. In this presentation, we report the detection of the $^{13}\text{C}17\text{O}$ J=3–2 emission line by ALMA observations (Band 7 (~0.9 mm), spatial resolution: ~0.3-0.4") in the disk around the Class I protostar V883 Ori, a typical young outbursting FU Ori type star. We found that the $^{13}\text{C}17\text{O}$ line emission has ring-like distribution within the disk surrounding the optically thick dust emission (radius $r \sim 40$ au) near the protostar, similar to the $^{12}\text{C}17\text{O}$ line emission line obtained in past observations ([4]). In addition, the $^{13}\text{C}17\text{O}$ line emission extends further to the outside of the disk ($r \sim 200$ au) compared with H_2O and CH_3OH line emission ($r \sim 80$ au). We derived the column density and the total gas surface distributions in the disk radial direction, and discuss the gas to dust mass ratio G/D, Toomre Q values (= Index of gravitational stability in the disk), and $^{12}\text{C}/^{13}\text{C}$ isotope ratio in the V883 Ori disk. We also conduct the comparative discussion of the $^{12}\text{C}/^{13}\text{C}$ ratio between CO (this study) and complex organic molecules (Yamato, Notsu et al. in prep.).

References

- [1] A. Miotello et al. 2017, A&A, 599, A113.
- [2] A. Booth et al. 2019, ApJL, 882, L31.
- [3] A. Booth & J. Ilee. 2020, MNRAS, 493, L108.
- [4] J. Tobin, 2023, Nature, 615, 227.